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ena are caused by a dark body rushing violently through a nebulous cloud seems to be too special to fit a uniform class of phenomena. Whether a stupendous volcanic outburst from one star, whatever its cause, or the collision of two dark bodies, could supply the observed facts, remains in doubt. The spectrum contains sufficient proof that the atmosphere of the star is disturbed in the extreme. W. W. CAMPBELL.

March 31, 1901.

NOVA PERSEI AND THE FAINT STARS NEAR IT.

The following measures were made for the purpose of locating the faint stars that are visible in the telescope close to the new star discovered by ANDERSON. All the stars that could be seen with the 36-inch telescope within 3′ of the *Nova* were measured, as well as a few that are more distant. As all of these stars are quite faint, the new star was also connected with the 9.0-magnitude star (Bonn A. G. Catal., 2948), which is about 6′ north of and preceding it.

The observing conditions on all the nights available for these measures were not very favorable, and the measures of distance were nearly all difficult. This was especially the case with the nearest star (a), which was at the limit of vision. The estimates of magnitude must be considered as relative rather than absolute. In addition to the estimates made on the nights of the measures, the following comparative estimate was made on March 29th: $o = 12.5$; $k = i = f = 13.0$; $l = m = 13\frac{1}{4}$; $j = c = 13.5$; $g = h = n = 14.0$; $b = 14\frac{3}{4}$; $e = 15.5$; $d = 16.0$; a, not visible.

The diagram facing page 65 shows the relative positions of these stars.

Nova and a.

Date.	Angle.	Distance.	Magnitude.
1901.184	348°.0	39″.39	—16.5
.203	348 .7	40 .28	—16.5
1901.19	348°.4	39″.82	—16.5

b.

Date.	Angle.	Distance.	Magnitude.
1901.184	107°.9	90″.86	—14.5
.239	108 .1	91 .88	—14.5
1901.21	108°.0	91″.37	—14.5

c.

1901.159	117°.2	102″.31	—13.5
.184	116 .6	102 .30	—13.5
.239	116 .7	103 .33	—13.5
1901.19	116°.8	102″.65	—13.5

e.

1901.203	142°.7	121″.34	—15.5

f.

1901.159	211°.4	123″.96	—13.0
.184	210 .9	125 .13	—13 0
.239	210 .9	124 .76	—13.0
1901.19	211°.1	124″.62	—13.0

i.

1901.159	62°.8	158″.33	—13.0
.184	63 .2	158 .21	—12.5
.203	63 .3	157 .95	—13.5
1901.18	63°.1	158″.16	—13.0

k.

1901.203	114°.6	168″.86	—13.0
.239	114 .5	170 .64	—13.5
1901.22	114°.5	169″.75	—13.2

n.

1901.184	352°.0	192″.75	—13.5
.239	352 .3	193 .39	14.0
1901.21	352°.2	193″.07	13.8

o.

1901.159	300°.3	236″.35	—13.0
.184	300 .4	236 .63	—12.0
1901.17	300°.4	236″.49	—12.5

p (= Bonn A.G. 2948).

1901.176	321°.6	367″.73	—9.0
.184	321 .7	367 .45	—9.0
1901.18	321°.6	367″.59	—9.0

e and *d.*

1901.203	306°.7	19″.50	—16.0

i and *g.*

1901.203	270°.4	30″.91	—13.5
.239	269 .6	31 .34	14.5
1901.22	270°.0	31″.12	—14.0

i and *h*.

1900.203	194°.5	45″.42	—13.5
.239	195 .4	45 .97	—14.0
1901.22	195°.0	45″.70	—13.8

o and *j*.

1901.159	114°.9	71″.71	—13.0
.184	114 .7	72 .78	—14.0
.239	114 .5	71 .85	—13.5
1901.19	114°.7	72″.11	—13.5

o and *l*.

1901.159	144°.2	61″.39	—13.5
.184	144 .2	61 .68	—13 0
.239	144 .3	61 .99	—13.0
1901.19	144°.2	61″.69	—13.2

k and *m*.

1901.203	100°.4	21″.39	—13.5
.239	100 .2	22 .01	—13.0
1901.22	100°.3	21″.70	—13.2

R. G. AITKEN.

THE VELOCITY OF GROOMBRIDGE 1830 IN THE LINE OF SIGHT.

The star No. 1830 in GROOMBRIDGE'S catalogue has been called the "Runaway Star," on account of its large proper motion. Its position in the heavens is changing at the rate of 7″.05 per year,—a speed sufficient to carry it over one degree in five hundred years. This was the maximum stellar motion known up to the year 1898, at which date an 8th-magnitude star in the southern hemisphere was found to have a proper motion of 8″.7 per annum.

Several determinations of the parallax of Groombridge 1830 have been made. The separate results differ widely in value, but they are in substantial agreement in placing the star at a great distance; it is probable that at least twenty stars with large proper motions are nearer than Gr. 1830. From a consideration of the merits of the individual determinations, NEWCOMB has adopted 0″.14 as its most probable parallax. Assuming this to be its true value, the component of the star's velocity at right angles